

## Previous Lecture

- Cerenkov effect
- Its application to SuperKamiokande
- Start atmospheric neutrinos...

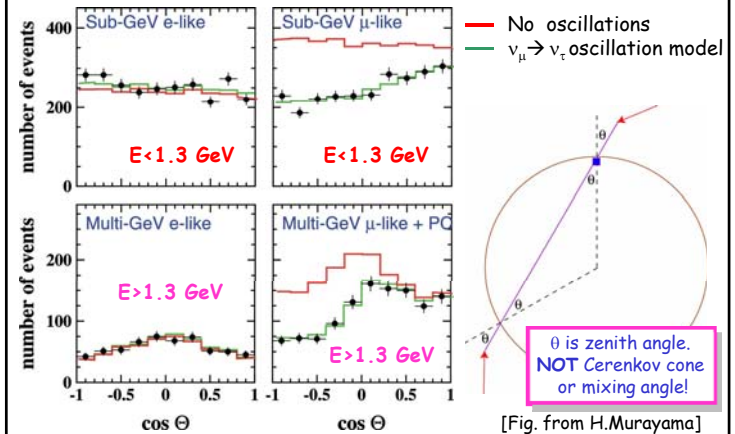
## Lecture Content

- Approx. lecture content
  1. PP intro
  2. PP intro.
  3.  $\nu$  props 1: strong/e.m./weak, no. neutrino generations
  4.  $\nu$  props 2: lepton no.,  $\nu$  existence
    - Examples of decay/production
  5. Neutrino mass
    - Fermi-Kurie plot
    - Phase space kinematics/4-momentum
  6. Parity and CP violation... (why so important in lepton sector?)
    - Wu et al.,  $^{60}\text{Co}$  experiment
  7. Detection & observation
    - Liquid, solid, bubble chamber
    - "Direct" methods (DONUT)
  8. Atmospheric neutrinos
    - Cerenkov detectors
    - SuperKamiokande experiment
  9. Atmospheric neutrino data and oscillations
    - Interpretation of atmospheric  $\nu$  data
    - Two-flavour neutrino oscillation formalism
  10. Solar neutrinos and SSM
    - SNO experiment and data
    - NDBD (NEMO, etc.)
  11. Implications for cosmology
    - Open vs. closed scenarios: various  $m$ , regions
    - $\nu$  as DM candidate?
    - Subject outlook (JPARC, MICE, Neutrino Factory, SK, SNO, KAMLAND, CHOOZ, MINOS, miniBOONE, JPARC,  $\nu\text{F}$ ).

## Today

- Atmospheric  $\nu$  data and their interpretation
- Neutrino oscillation concept
- 2 flavour neutrino flavour oscillation formalism
- Generalisation to 3 flavours: MNS (Maki-Nakagawa-Sakata) mixing

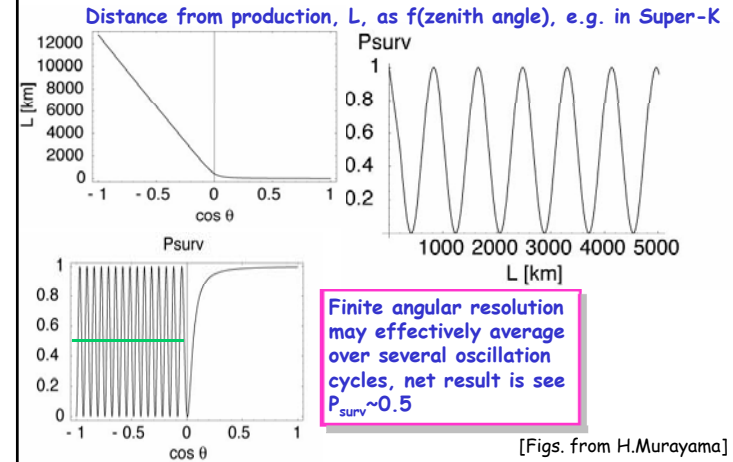
## Super-Kamiokande results II



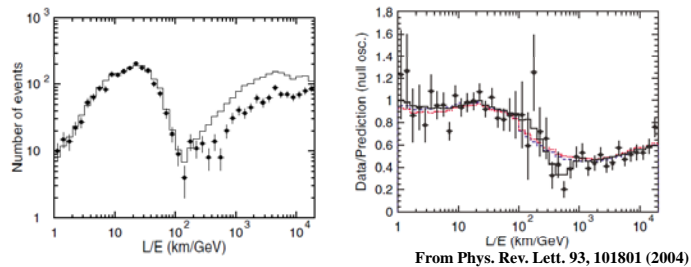
## Super-Kamiokande results I

- Super-K atmospheric  $\nu$  data, solid evidence for oscillations (1998)
- Predicted  $N(\mu)/N(e)$  ratio  $\cong 2$ 
  - ▶  $\mu$  from  $\pi^- \rightarrow \mu^- + \bar{\nu}_\mu$  and  $\mu^- \rightarrow e^- + \nu_\mu + \bar{\nu}_e$
  - ▶  $e$  only from  $\mu^-$  decay
- Measured  $N(\mu)/N(e)$  / predicted  $N(\mu)/N(e) = 0.64 \pm 0.05$ 
  - ▶ Prediction includes simulation of resolution, detector thresholds, etc.
- $N(\mu)/N(e)$  increases for higher energy  $\mu$  (time dilation by Lorentz  $\gamma$  factor)
  - ▶ More time dilation  $\Rightarrow$  fewer  $\mu$  decay in atmosphere:  $R = \gamma c \tau = (E/m) c \tau$
  - ▶ Remember  $\mu$  from  $\pi^- \rightarrow \mu^- + \bar{\nu}_\mu$  and  $\mu^- \rightarrow e^- + \nu_\mu + \bar{\nu}_e$
  - ▶  $e$  only from  $\mu^-$  decay
- Clear asymmetry in upward/downward  $\mu$  events
  - ▶ Shows no. of  $\nu_\mu$  depends on distance travelled between production (in atmosphere) and detector
  - ▶ No such effect for electrons
  - ▶ Common interpretation:  $\nu_\mu \rightarrow \nu_\tau$  oscillation

## "Survival" Probability



## Superkamiokande data



- Compare (r.h.s.) best-fit expectations for
  - ▶  $\nu$  decay (dashed) - disfavoured at  $3.4 \sigma$  significance
  - ▶  $\nu$  decoherence (dotted) - disfavoured  $3.8 \sigma$  significance
  - ▶  $\nu$  oscillations (solid) - preferred hypothesis